

The Interplay of Plasma Turbulence and Magnetic Reconnection in Producing Nonthermal Particles

Luca Comisso and Lorenzo Sironi

Department of Astronomy, Columbia University
Columbia Astrophysics Laboratory, Columbia University

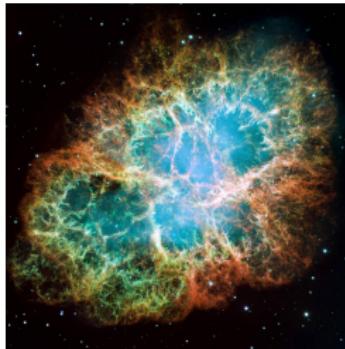
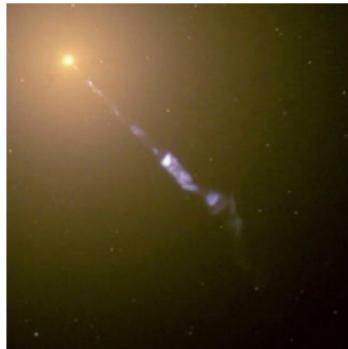
Heliophysics Seminar
PPPL, Princeton, April 12 2019



COLUMBIA UNIVERSITY
IN THE CITY OF NEW YORK

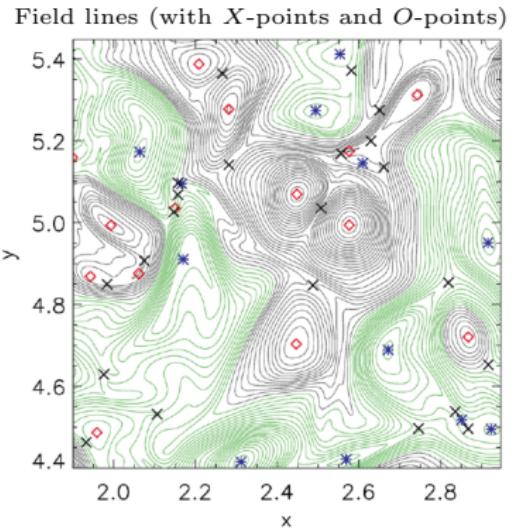
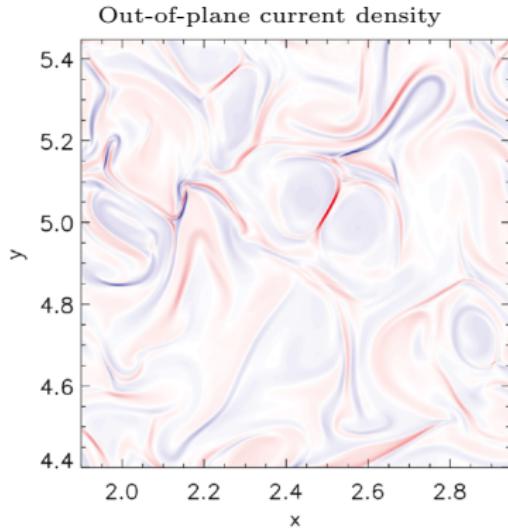
Outline

- ▶ Is turbulence in magnetically-dominated (i.e., relativistic) plasmas an efficient source of nonthermal particles?
- ▶ If so, how does the nonthermal spectrum depend on the system parameters?
- ▶ Mechanism of particle acceleration? Interplay with magnetic reconnection?



Nonrelativistic Turbulence (plus Reconnection)

- ▶ Plasma turbulence produces sheets of strong current density \Rightarrow natural sites of magnetic reconnection

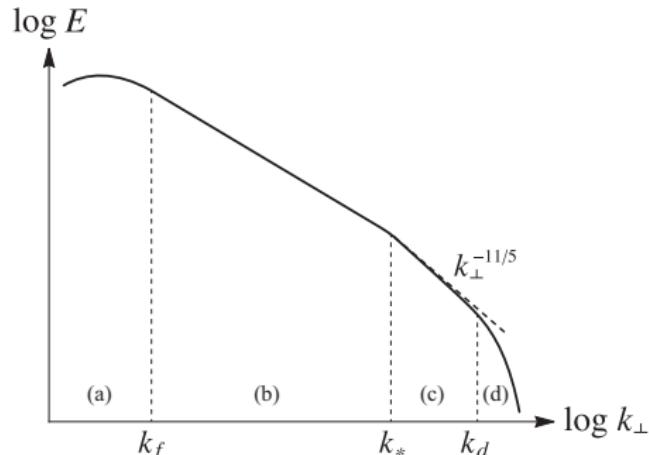


(Servidio *et al.* 2009/10/11)

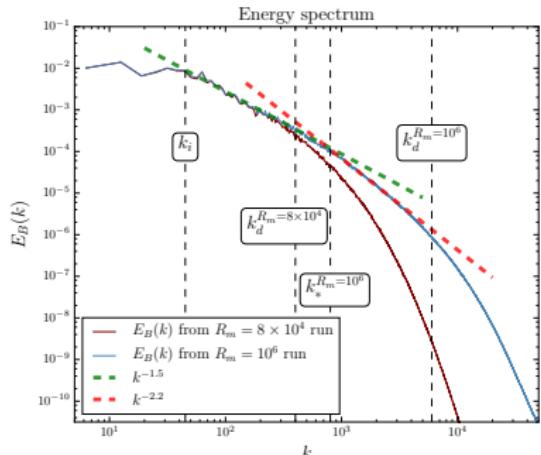
(also Matthaeus & Lamkin 1986, Politano *et al.* 1989, Biskamp 2003,
Mininni *et al.* 2006, Wan *et al.* 2013, Zhdankin *et al.* 2013)

Nonrelativistic Turbulence (plus Reconnection)

- Magnetic reconnection leads to a steepening of the energy spectrum. What about particle acceleration?



(Comisso *et al.* ApJ 2018)



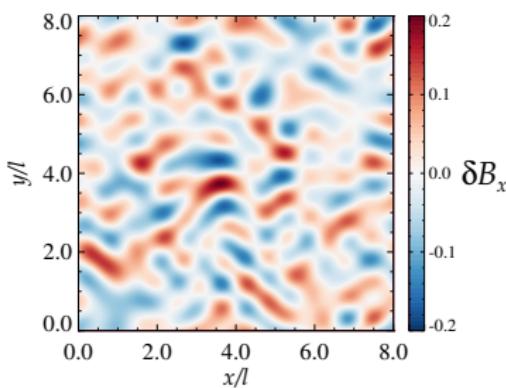
(Dong *et al.* PRL 2018)

(also Carbone *et al.* 1990, Carbone 1995, Cerri & Califano 2017, Franci *et al.* 2017, Mallet *et al.* 2017, Loureiro & Boldyrev 2017, Walker *et al.* 2018, ..)

Tools and basic setup

- ▶ Fully-kinetic treatment:
 - ⇒ we solve the coupled Vlasov-Maxwell system of equations through the PIC method
- ▶ PIC code TRISTAN-MP (Spitkovsky 2005)
- ▶ Decaying turbulence in relativistic pair plasmas
- ▶ 2D and 3D numerical simulations

Tools and basic setup (Relativistic “Servidio Problem”)



- ▶ mean magnetic field $\langle \mathbf{B} \rangle = B_0 \hat{\mathbf{z}}$
- ▶ turbulence develops from uncorrelated magnetic fluctuations δB_x and δB_y in Fourier harmonics
- ▶ energy-carrying scale: $l = 2\pi/k_f$ (k_f is the wavenumber where the energy spectrum peaks)

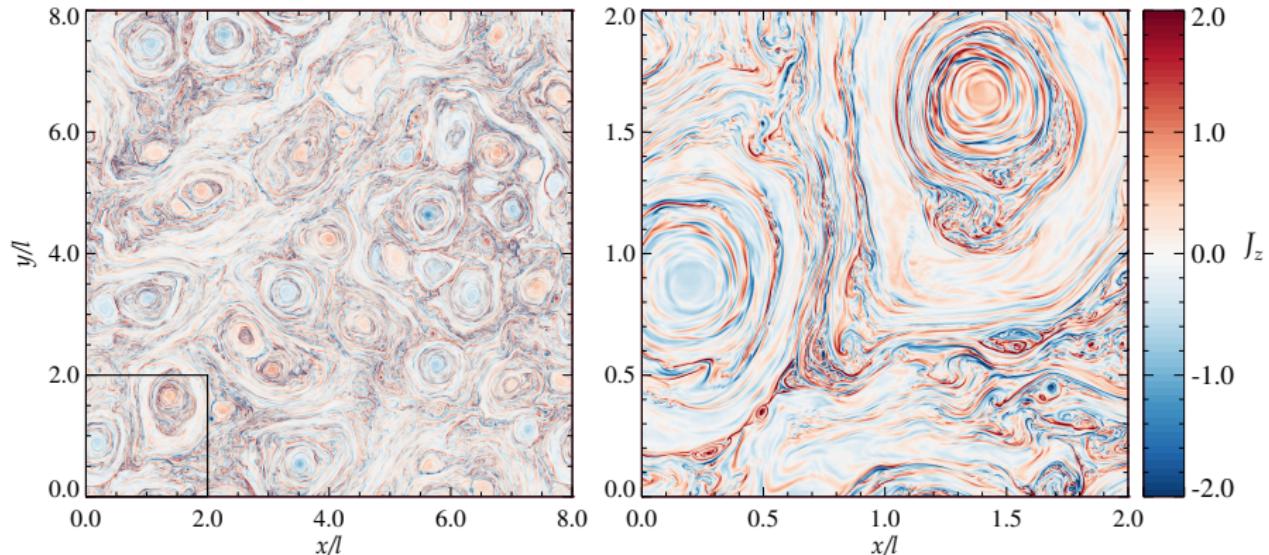
- ▶ Simulations in the magnetically-dominated regime

$$\sigma_0 = \frac{\delta B_{\text{rms}0}^2}{4\pi w_0} \gg 1, \quad \frac{1}{16} \leq \frac{\delta B_{\text{rms}0}^2}{B_0^2} \leq 16, \quad \theta = \frac{k_B T}{m_e c^2} \sim 1$$

with $w_0 = n m_e c^2 + n k_B T [\hat{\gamma}/(\hat{\gamma} - 1)]$

Fully-developed turbulence state

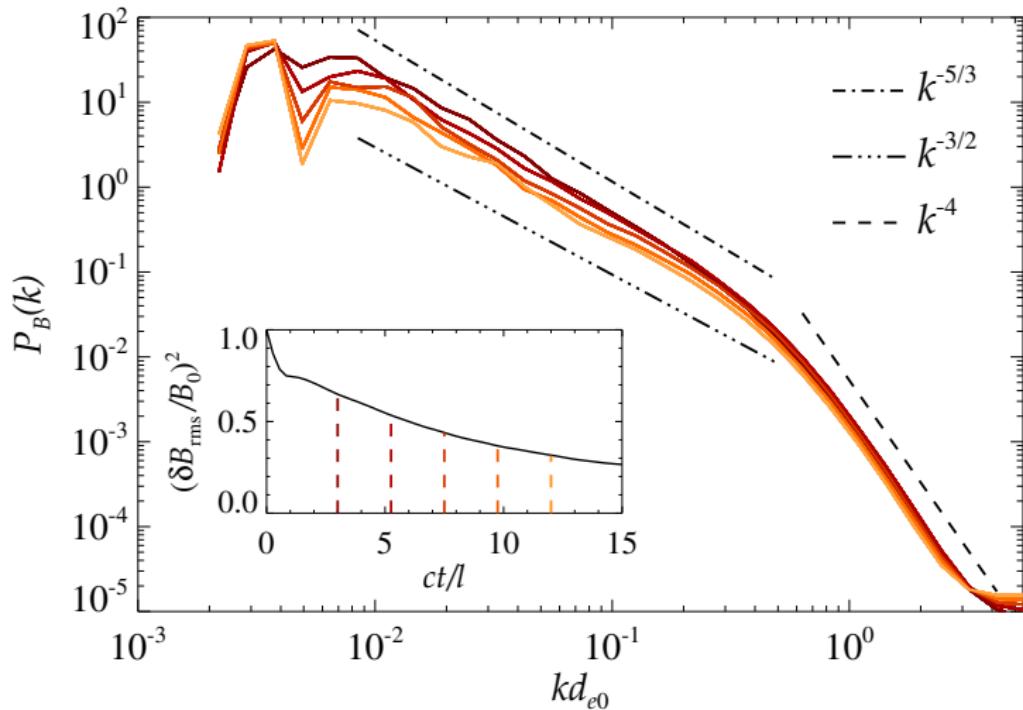
2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 3280$



- ▶ Copious presence of current sheets, plasmoids, and vortices

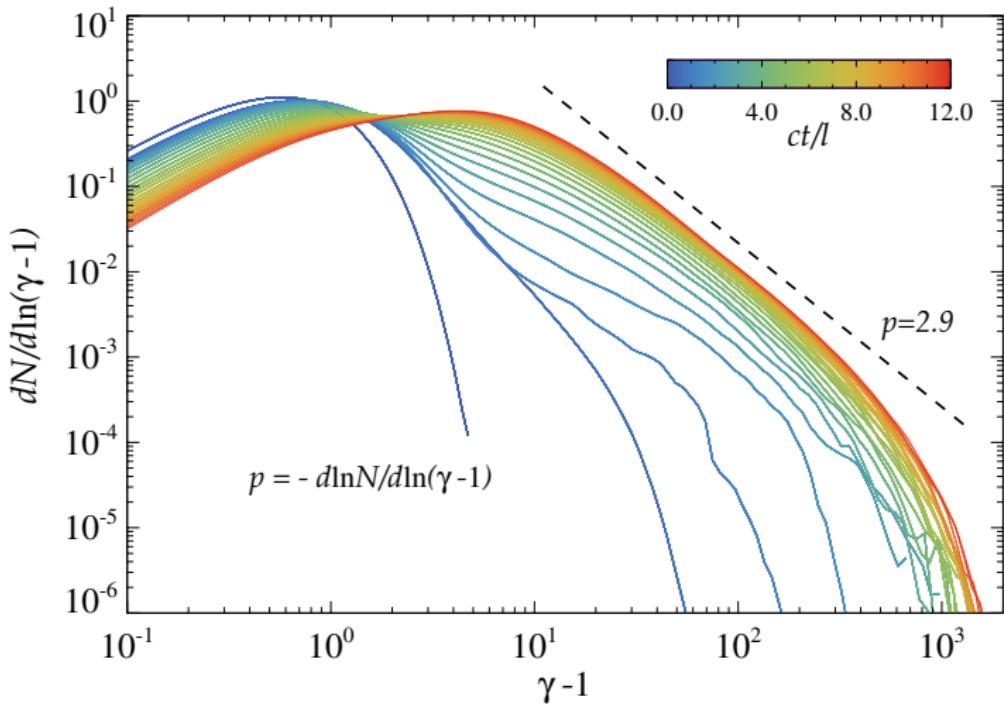
Fully-developed turbulence state

2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 3280$



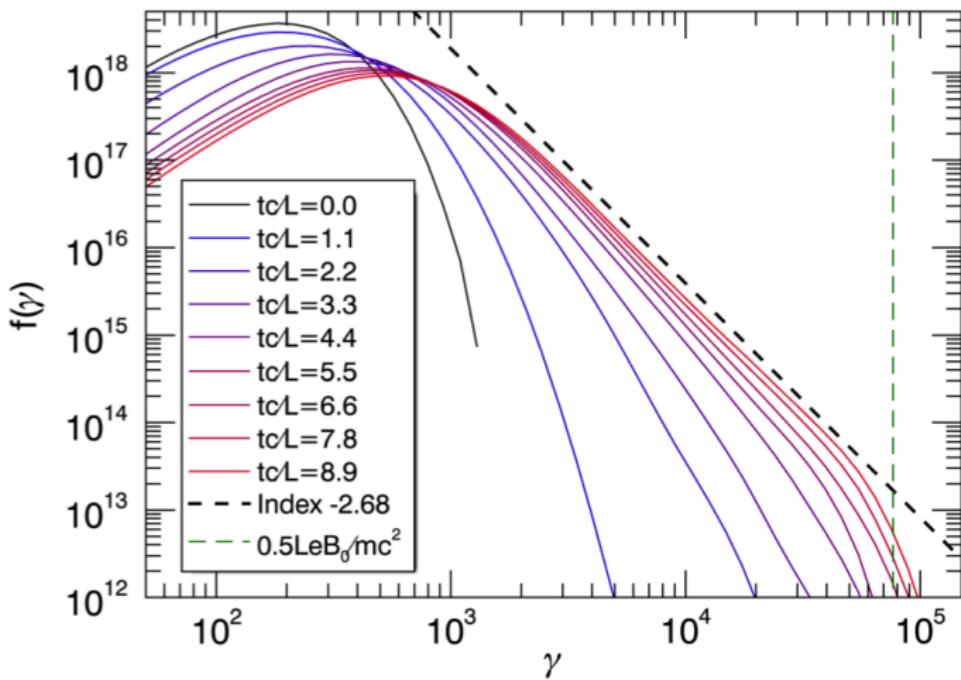
Particle Spectrum

2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 3280$



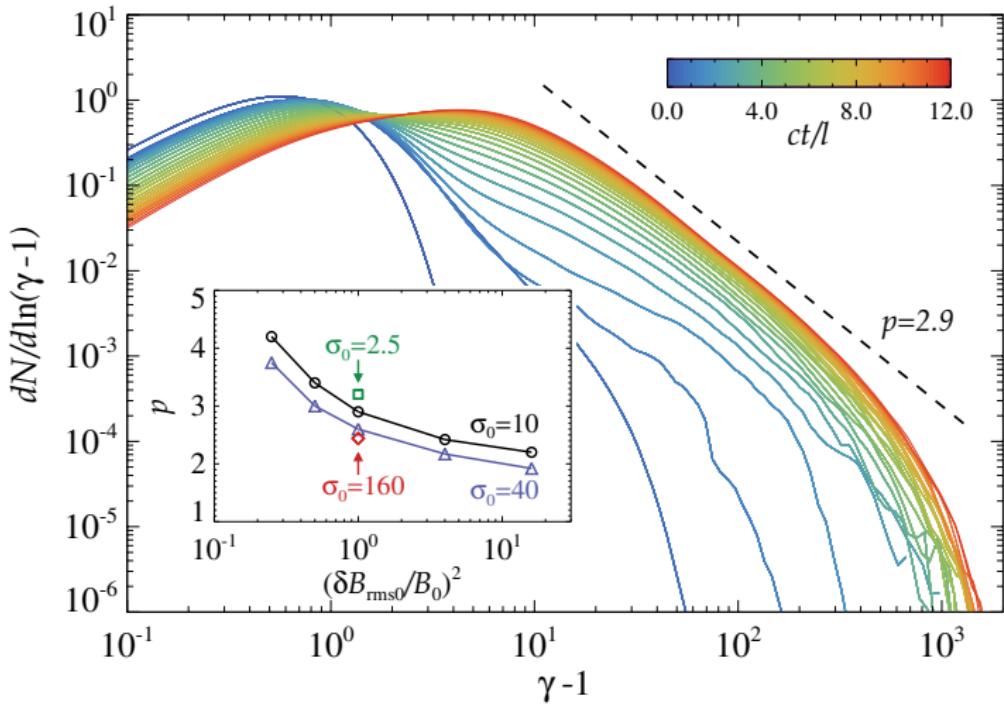
Particle Spectrum

Similar power-law particle energy distributions also in simulations by Zhdankin *et al.* 2017, 2018.



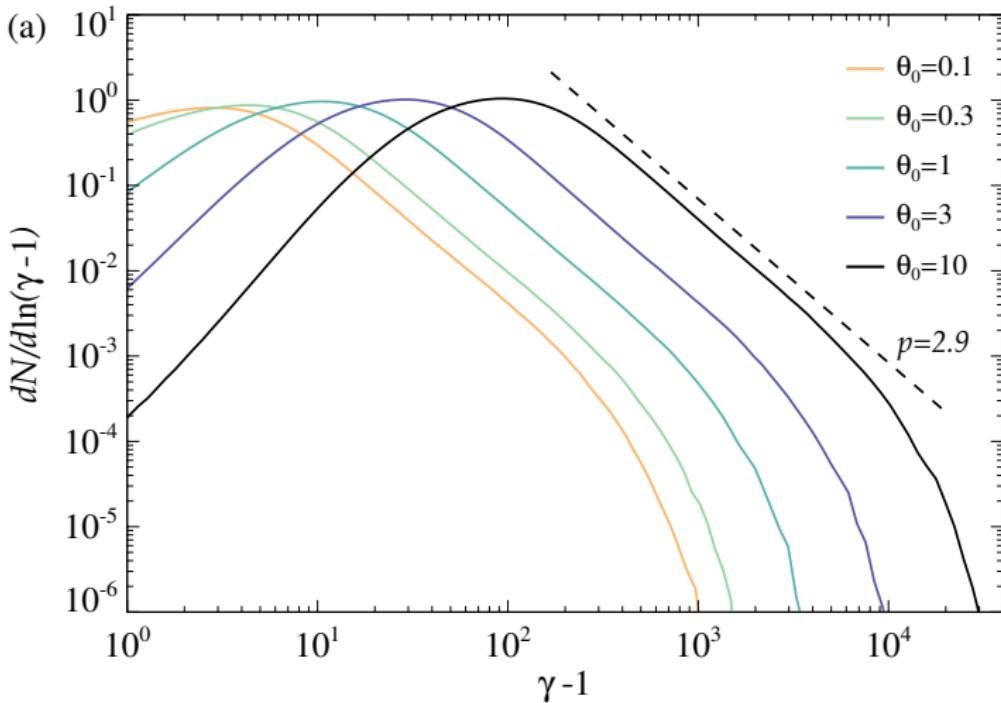
Particle Spectrum

2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 3280$



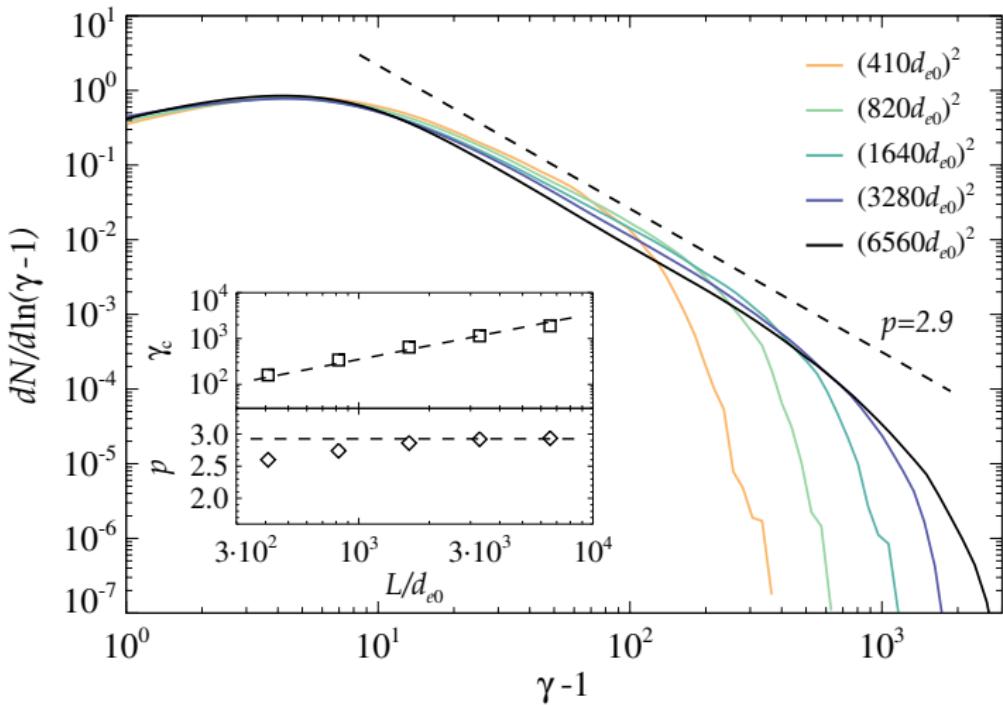
Particle Spectrum: initial temperature scan

2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 3280$



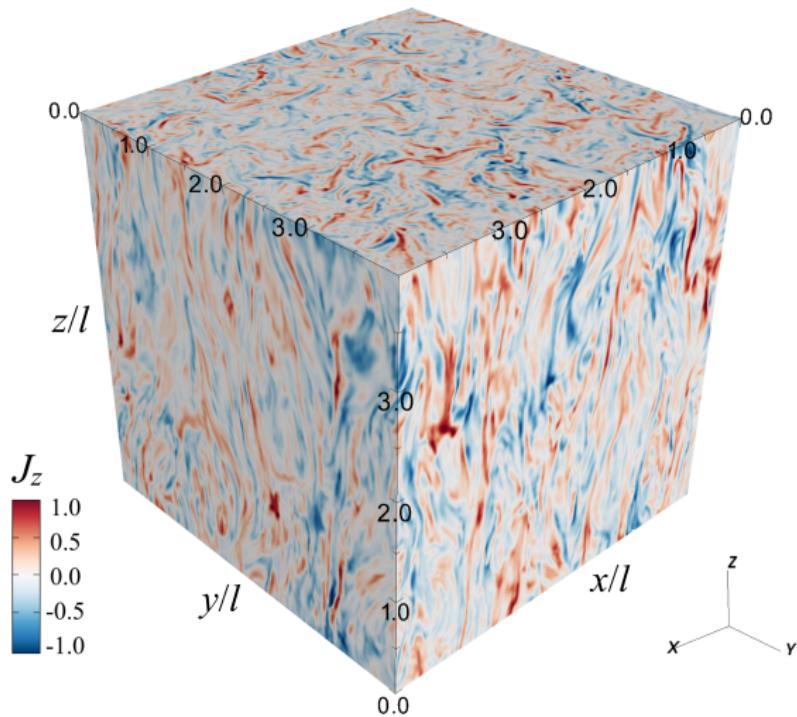
Particle Spectrum: system-size scan

2D simulations with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$

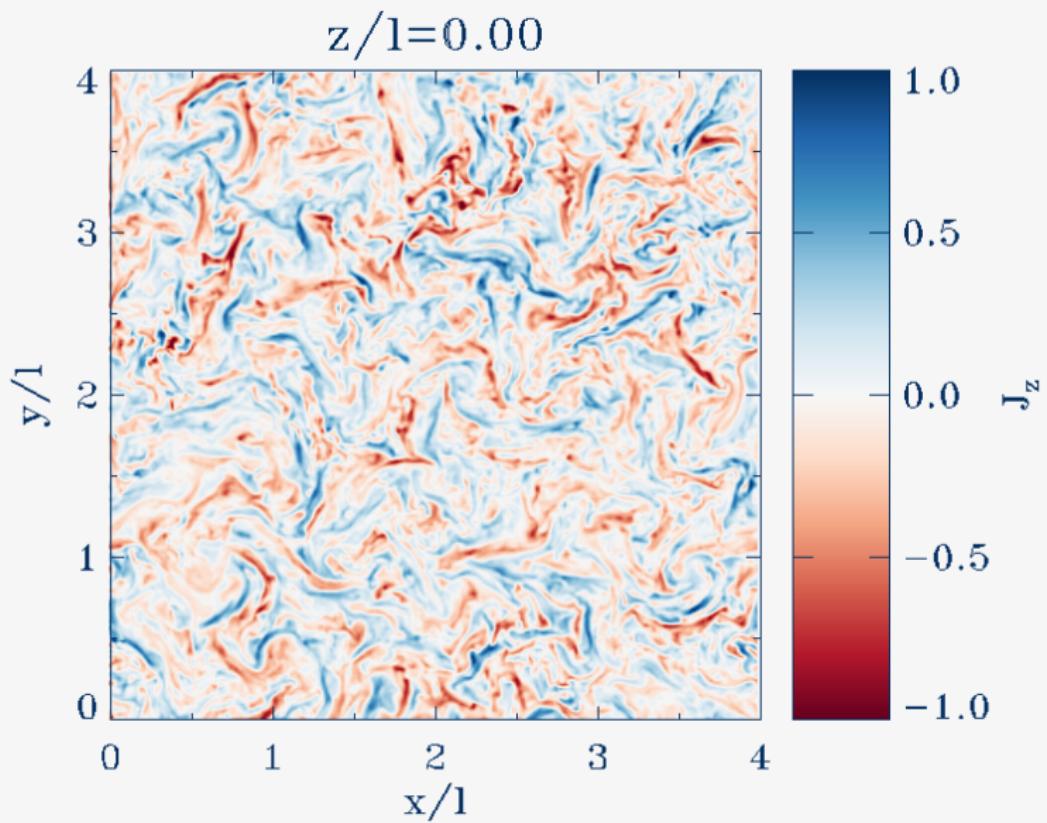


Current sheets in 3D

3D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 820$

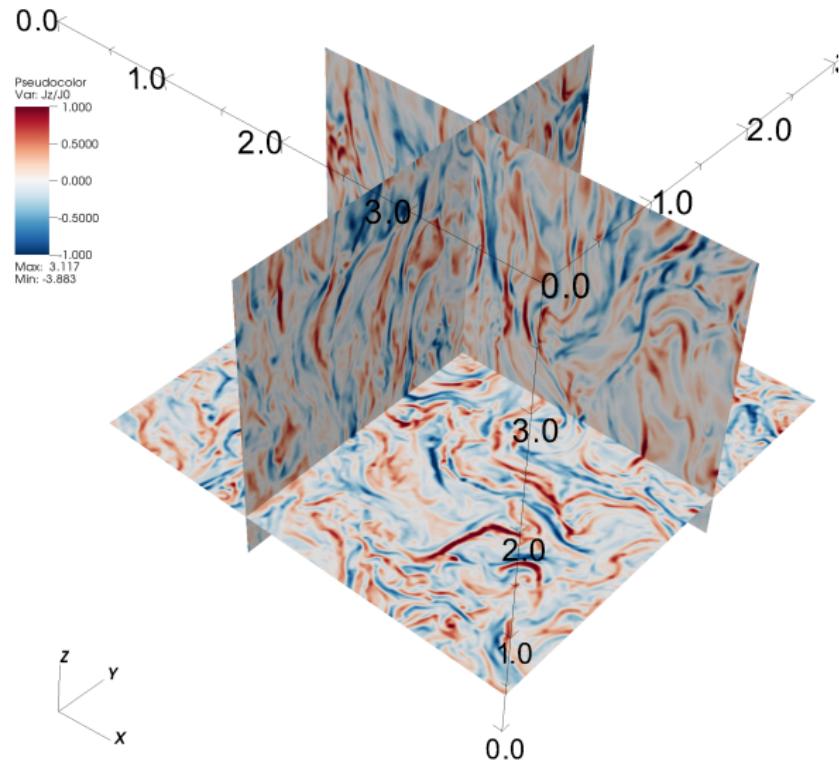


Current sheets in 3D: slices of J_z



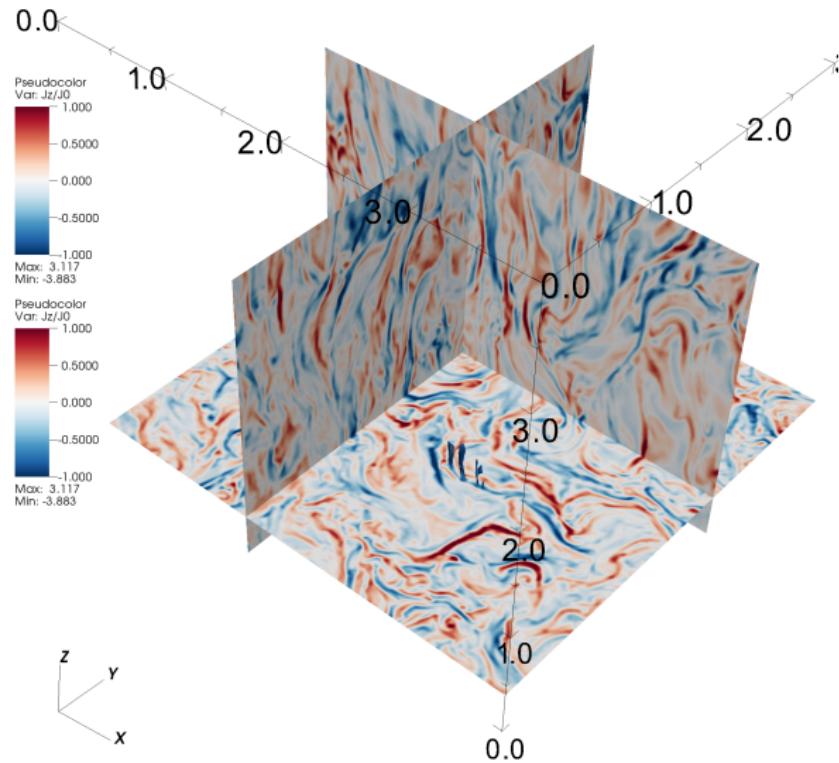
Current sheets in 3D: plasmoids/flux ropes

3D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 820$



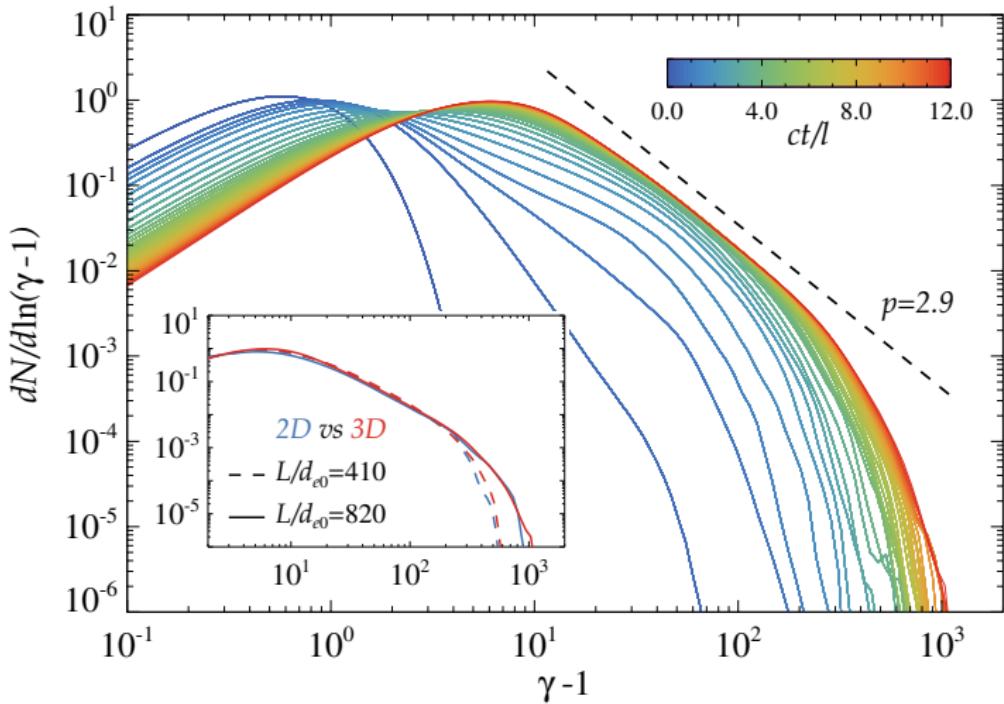
Current sheets in 3D: plasmoids/flux ropes

3D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 820$

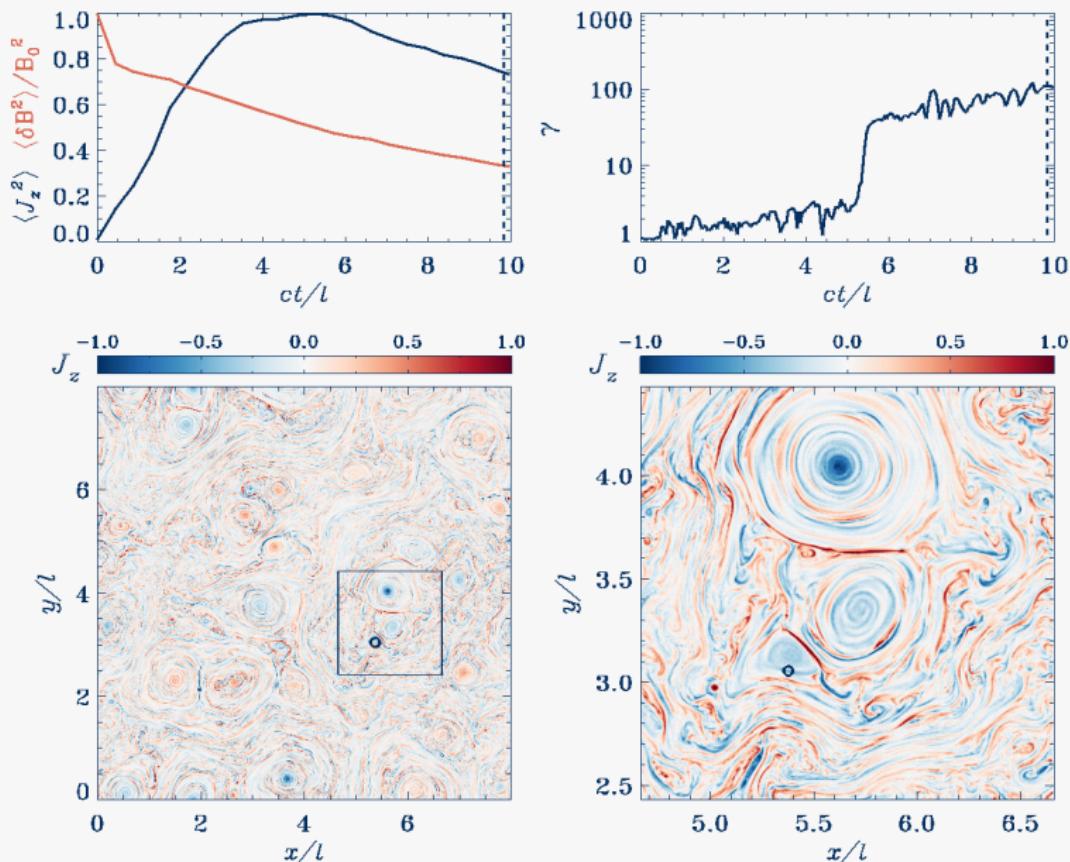


Particle Spectrum

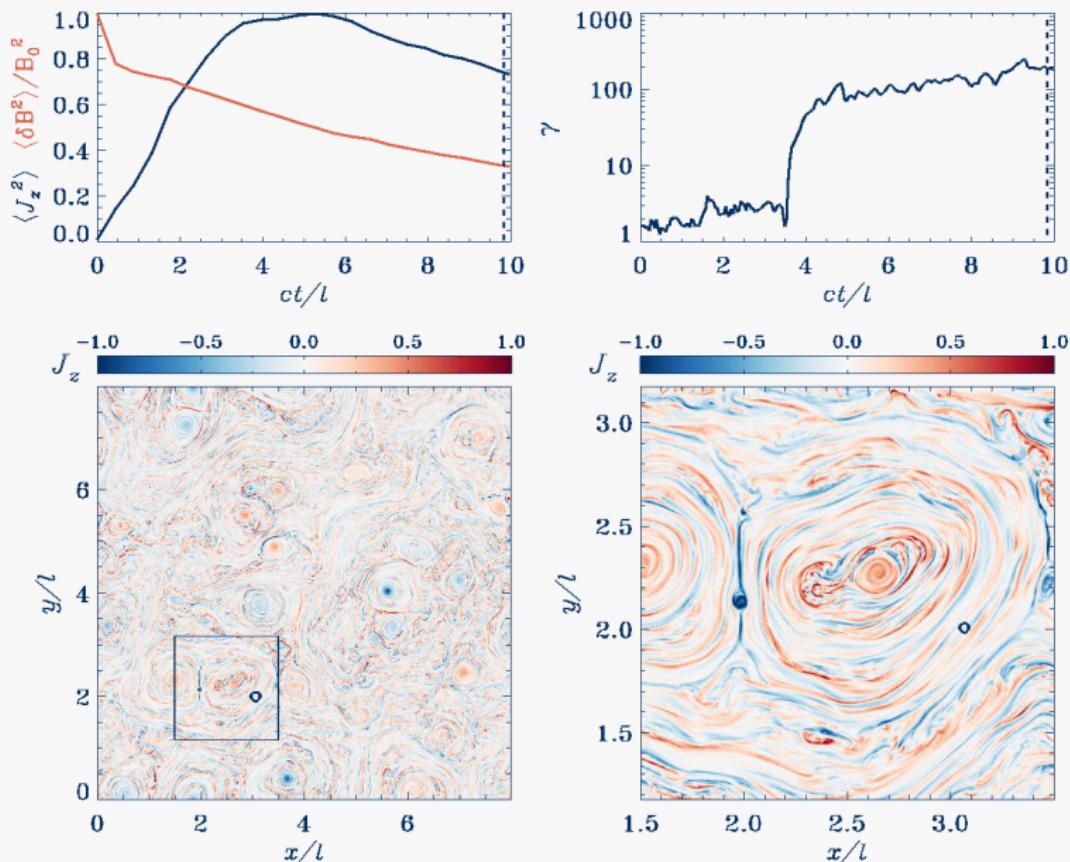
3D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 820$



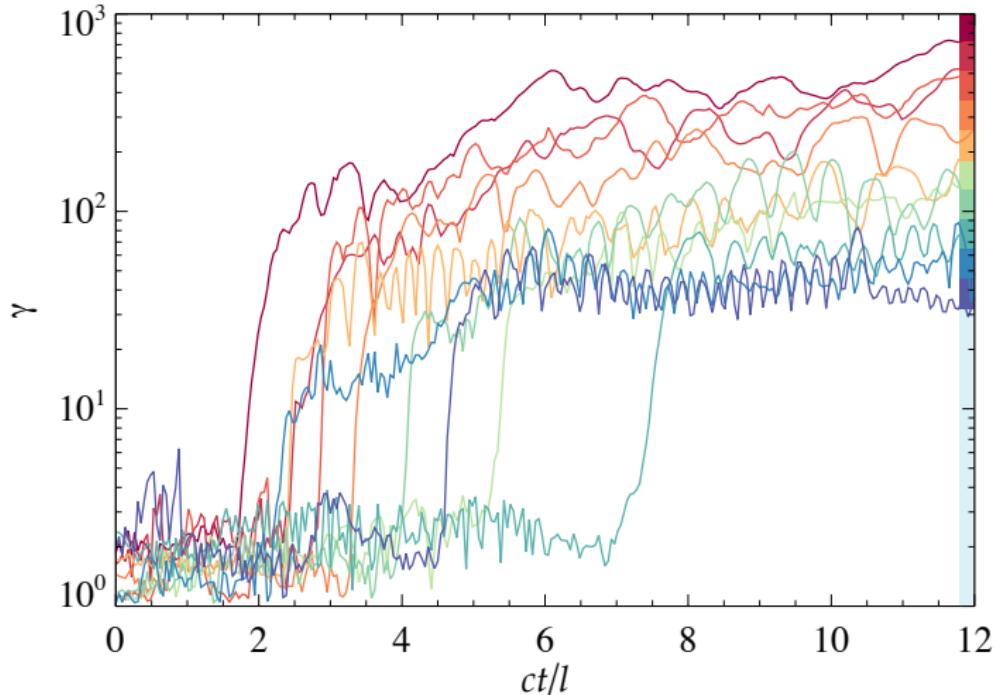
How are particles accelerated?



How are particles accelerated?

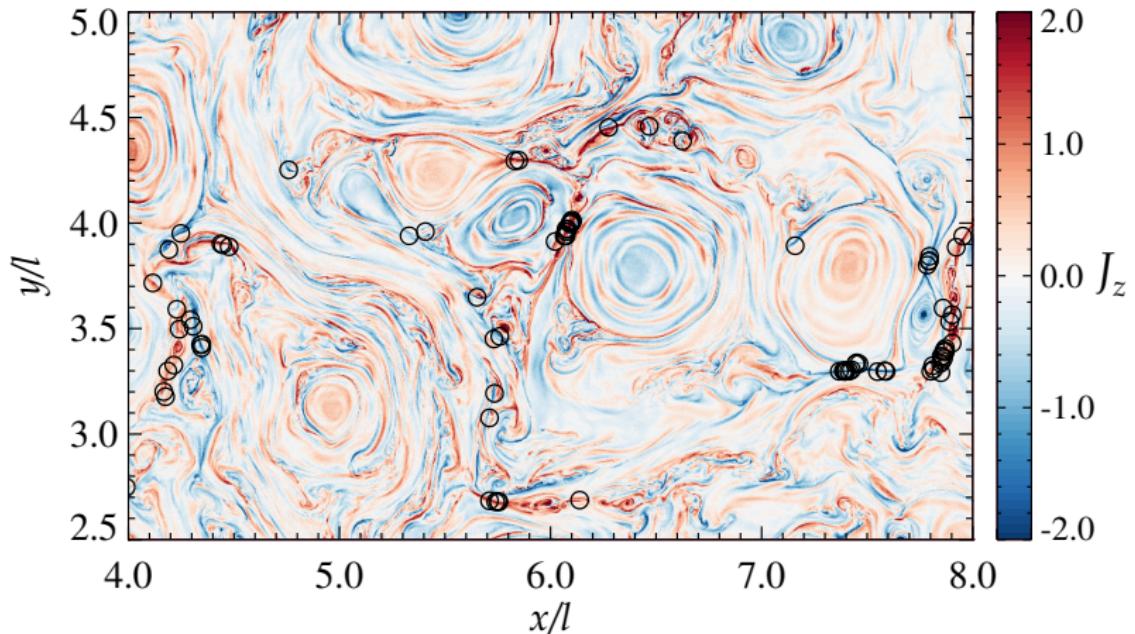


Particle Injection



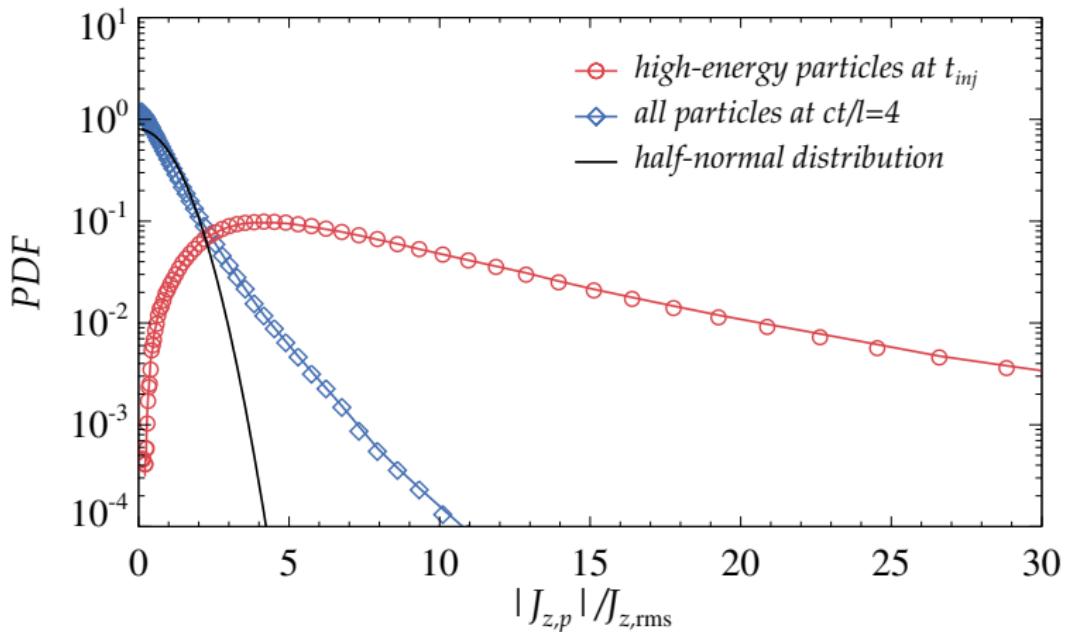
- ▶ $\sim 97\%$ of the particles belonging to the nonthermal tail experience a sudden energy jump from the thermal pool

Particle Injection



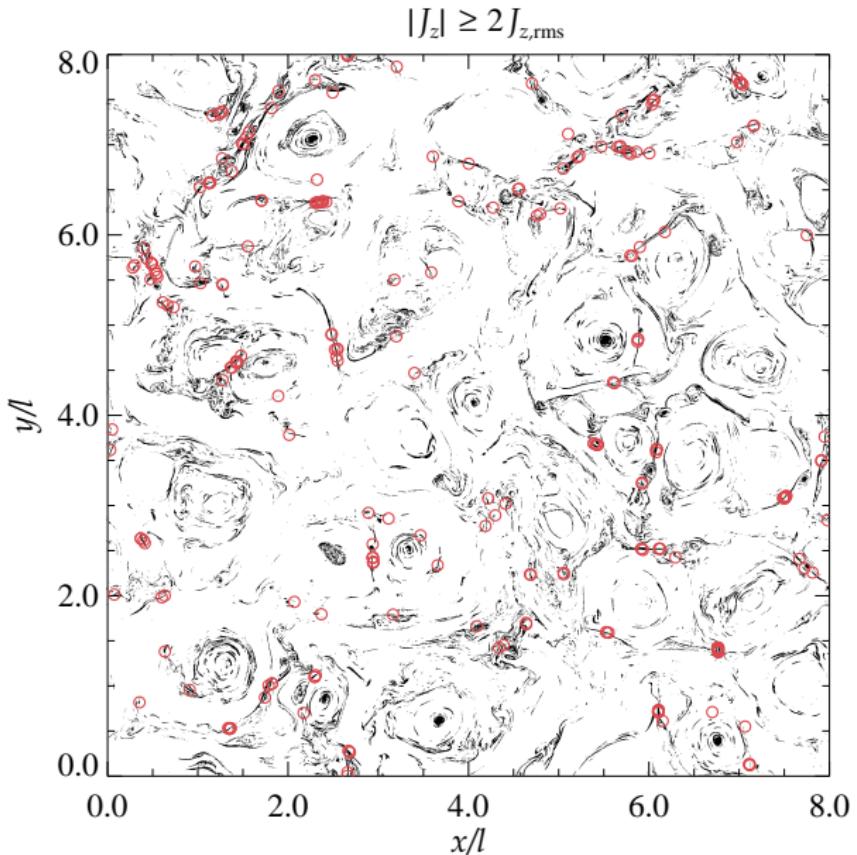
- ▶ Particle injection at reconnecting current sheets

Particle Injection

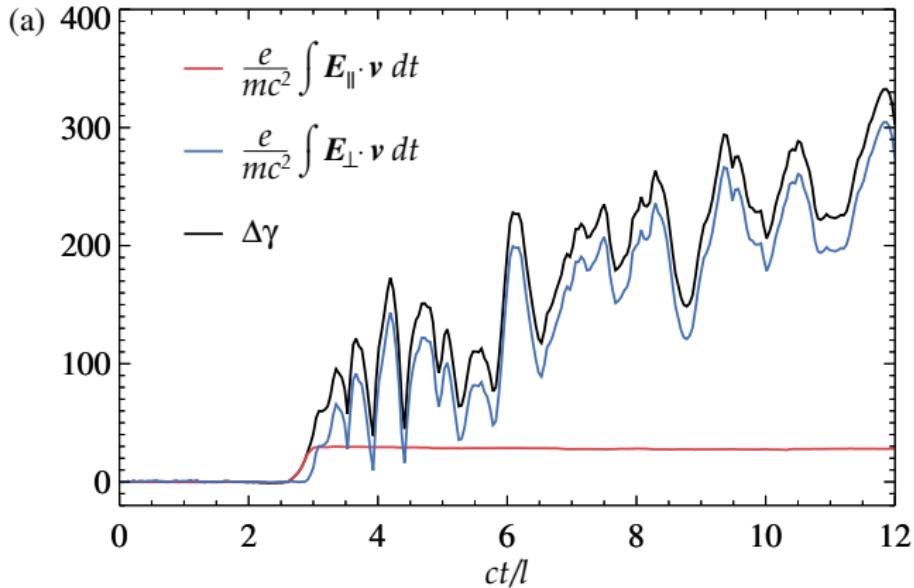


- ▶ $\sim 95\%$ of the high-energy particles reside at injection at $|J_{z,p}| \geq 2 J_{z,\text{rms}}$ (current sheets)

Particle Injection

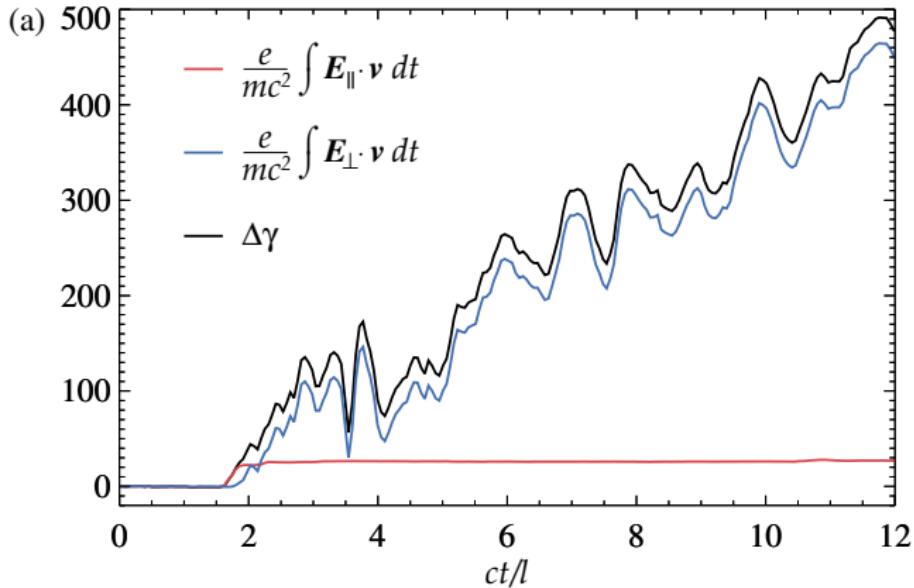


Particle energization



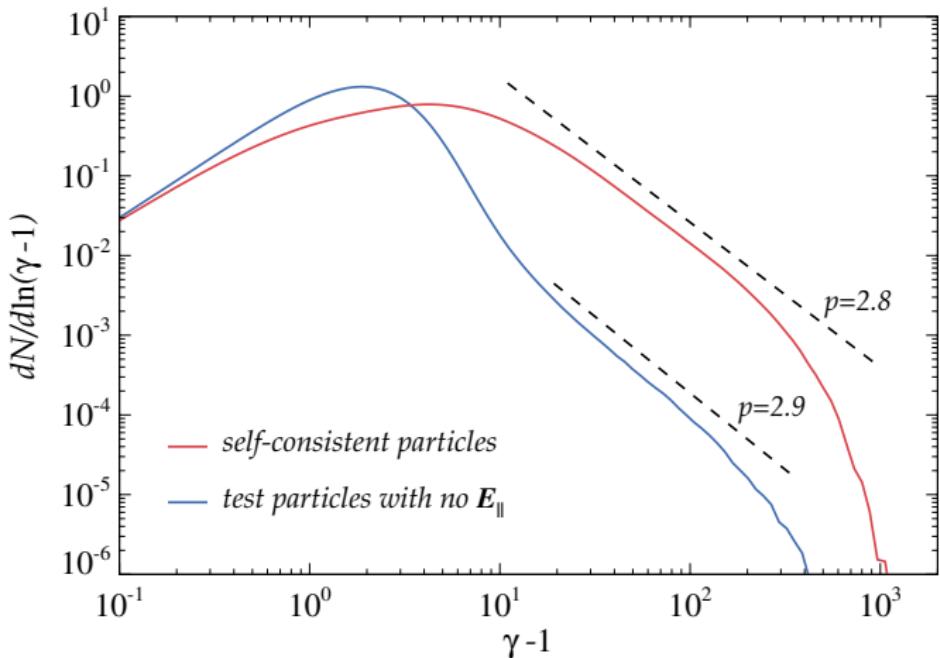
- ▶ Typical particle energization history (2D)

Particle energization



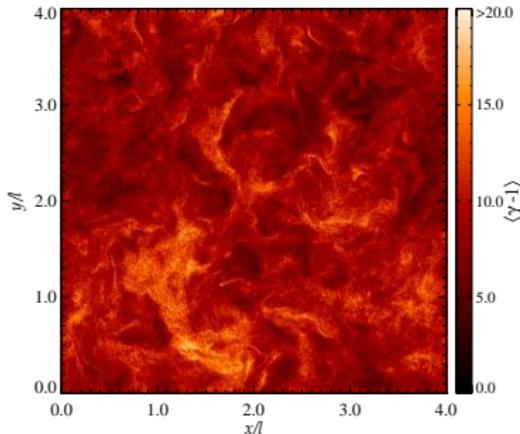
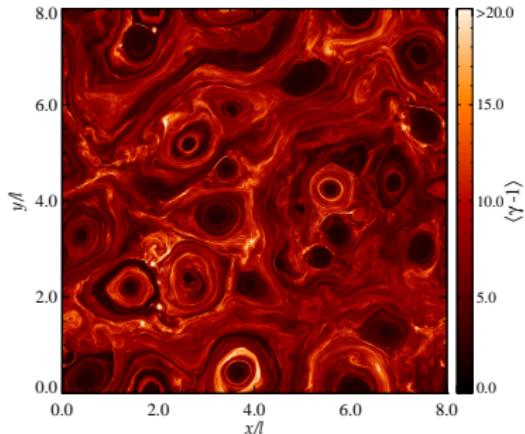
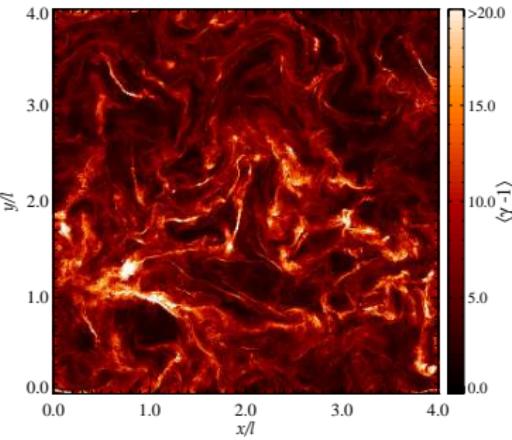
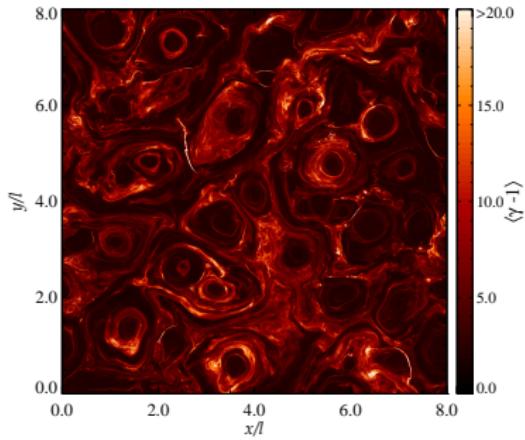
- ▶ Typical particle energization history (3D)

Test particle spectrum

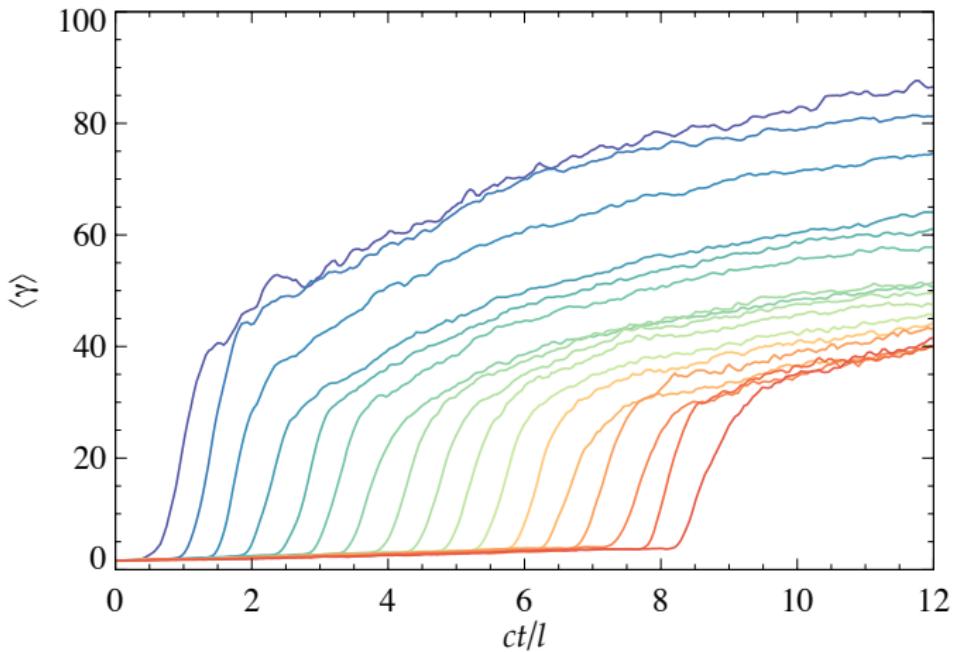


- ▶ For test particles that do not feel E_{\parallel} , there is a dramatic drop of particles in the nonthermal tail (while p is similar)

Particle mixing

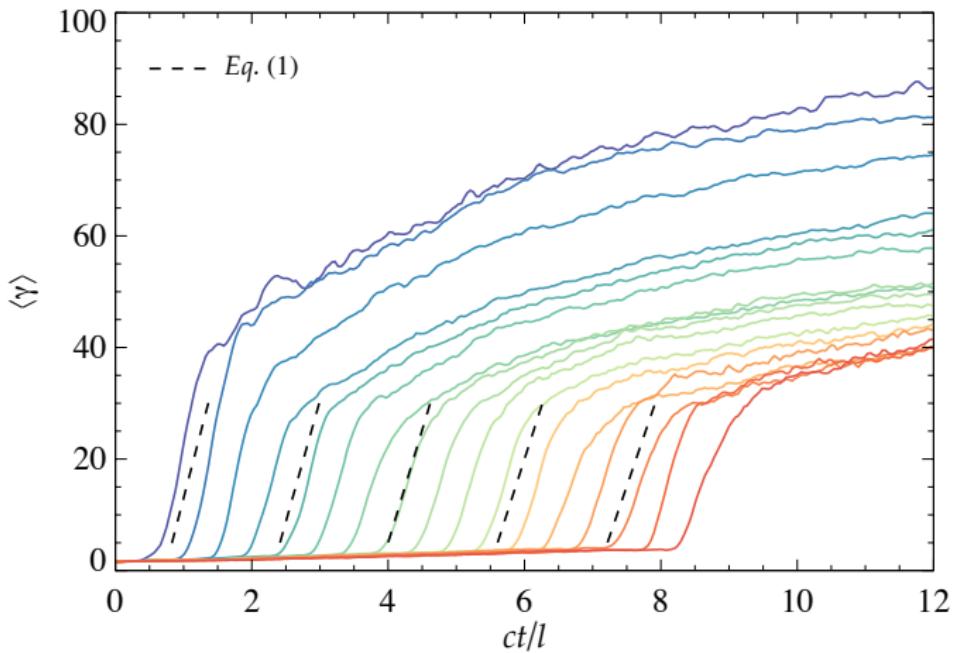


Energization mechanisms



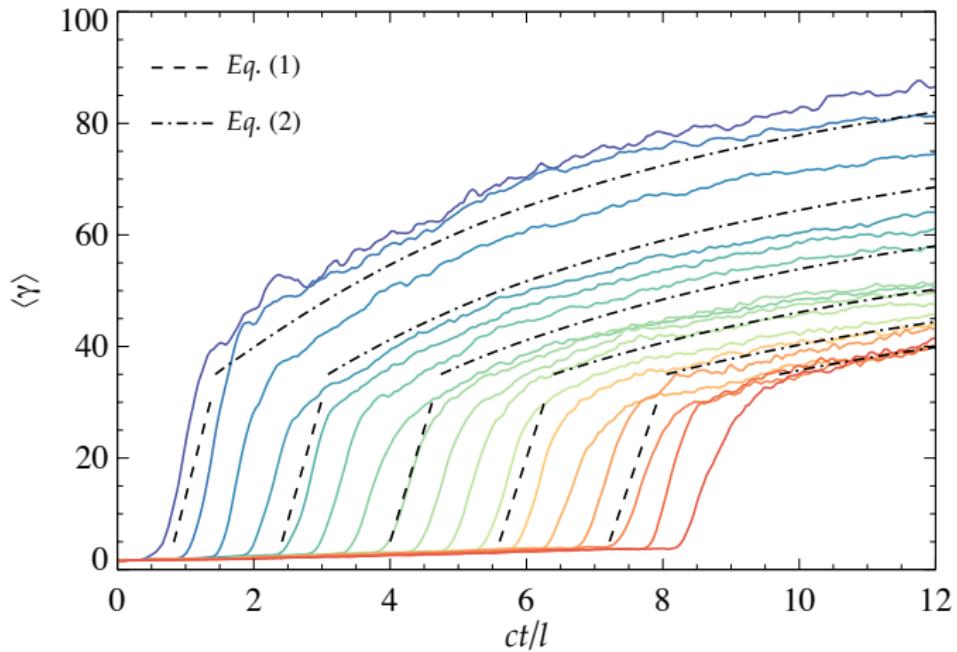
- ▶ Particles energization occurs in a two-stage process

Energization mechanisms: injection



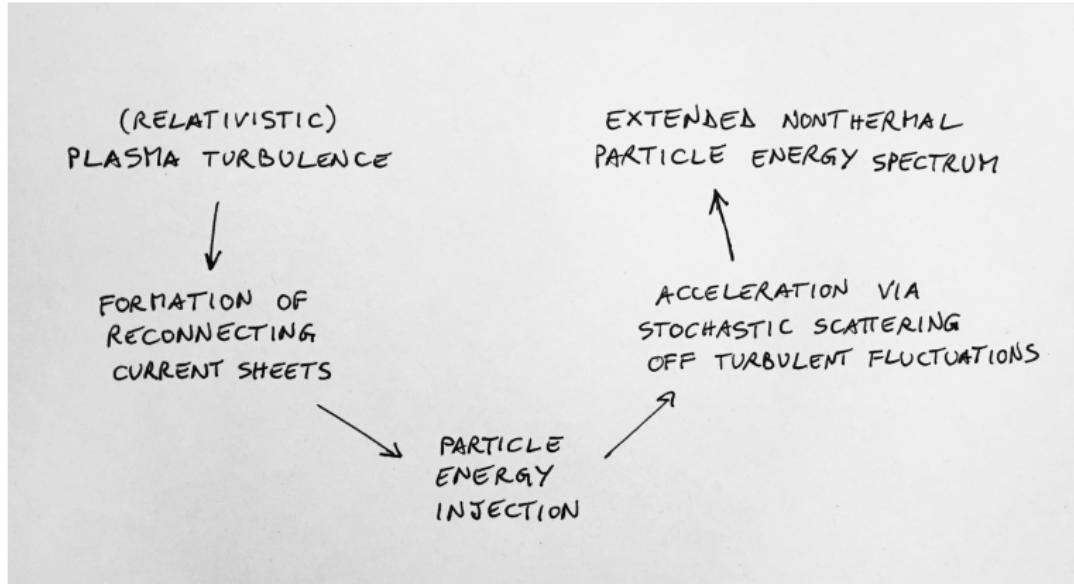
$$(1) \quad \frac{d\langle \gamma \rangle}{dt} = \frac{e}{mc} \beta_R \delta B_{\text{rms}}$$

Energization mechanisms: stochastic acceleration



$$(2) \quad \frac{d\langle\gamma\rangle}{dt} = \frac{1}{\gamma^2} \frac{\partial}{\partial\gamma} [\gamma^2 D_p] , \quad D_p = \frac{1}{3\kappa} \frac{\delta V_{\text{rms}}^2}{c^2} \frac{\delta B_{\text{rms}}^2}{B_0^2} \gamma^2 \omega_c$$

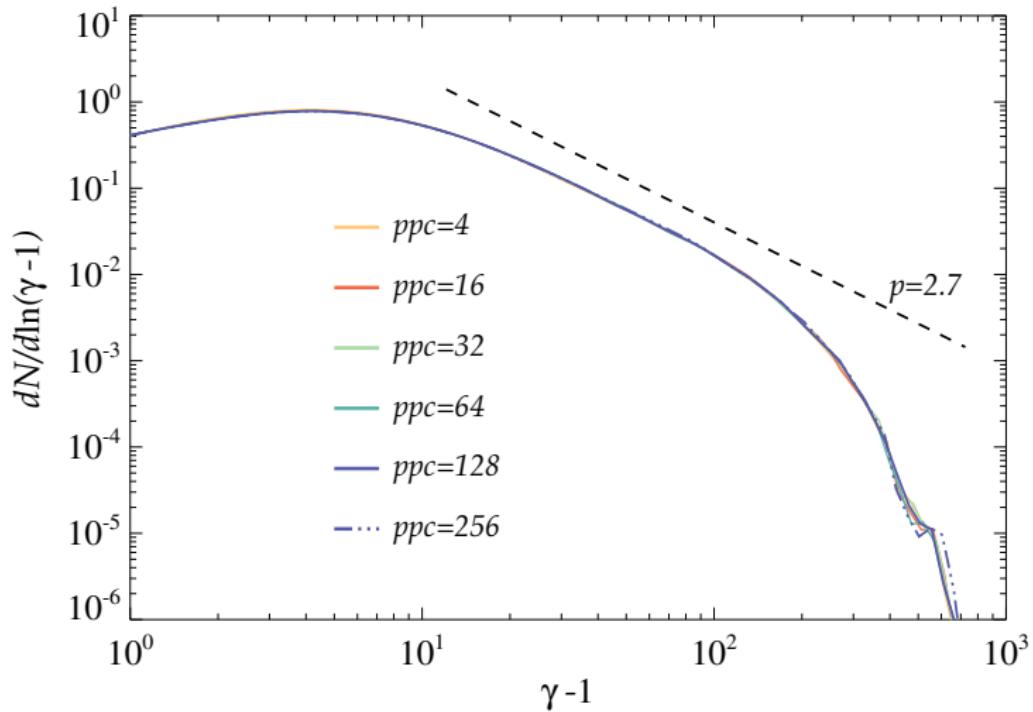
Summary



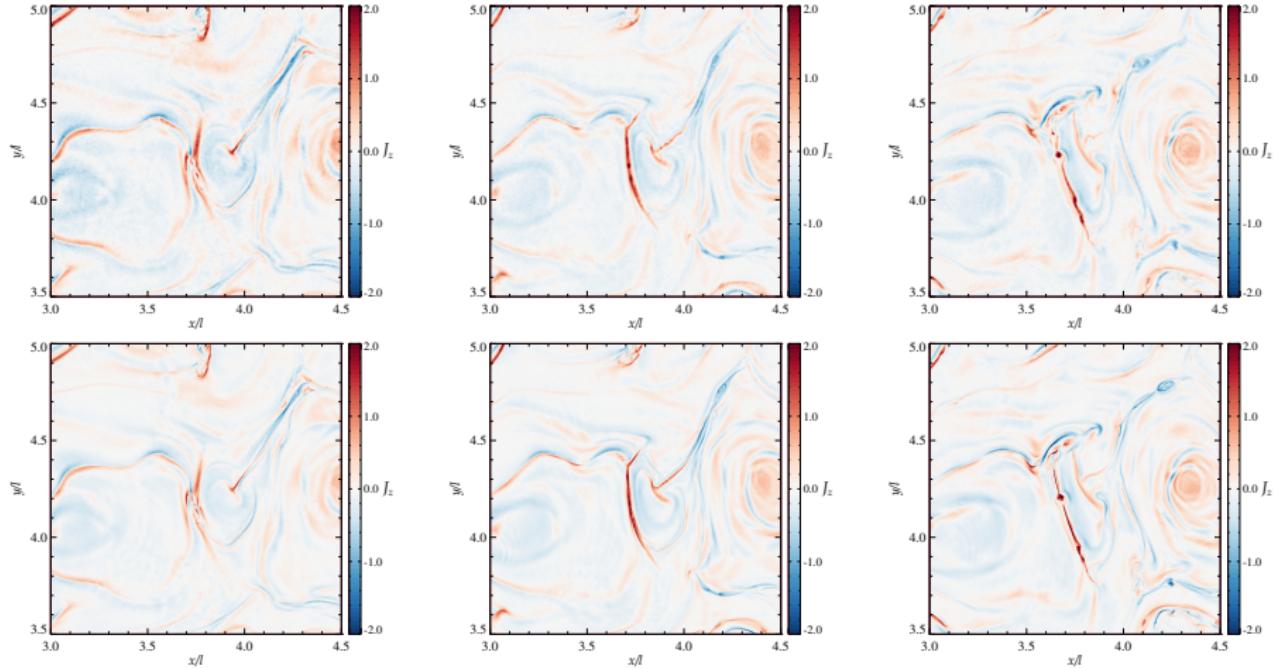
More details in Comisso & Sironi, PRL 121, 255101 (2018)

Particle Spectrum: particles per cell scan

2D simulation with $\sigma_0 = 10$, $\delta B_{\text{rms}0}/B_0 = 1$, $L/d_{e0} = 820$



Plasmoid formation: different spatial resolutions



From two 2D simulations with 3 cells per d_{e0} (top row) and 10 cells per d_{e0} (bottom row).

Movie - time evolution of J_z

